

ENERGY SPECTRUM OF EXTRAGALACTIC GAMMA-RAY SOURCES

R.J. Protheroe

Department of Physics, University of Adelaide
Adelaide, South Australia 5001

1. **INTRODUCTION.** Soon after the discovery of the cosmic microwave background radiation it was pointed out that γ -rays above 10^{14} eV would be attenuated by photon-photon pair production interactions with photons of the cosmic background (Jelley, 1966; Gould and Schreder, 1966, 1967a). Other processes, e.g. double pair production (Brown *et al.*, 1973), can also be important at extremely high energies. Later, it was realised that if the intergalactic magnetic field was low then γ -rays above 10^{14} eV might still reach Earth from extragalactic distances through an electron-photon cascade involving pair production and inverse Compton (IC) interactions (Bonnetto, 1971; Bonnetto and Lucchin, 1971; Allcock and Wdowczyk, 1972; Wdowczyk *et al.*, 1972; Stecker, 1973; Gould and Rephaeli, 1978). In many of these calculations, only the highest energy electron or photon produced in an interaction was considered in order to reduce the problem from a cascade calculation to a simpler one of calculating the energy loss rate of the "high energy particle".

Recently, Clay *et al.* (1984) reported an excess of EAS above $\sim 10^{15}$ eV from the direction of Cen A. At first sight, it seems unlikely that the excess could be due to γ -rays as Cen A is about 6 Mpc away, more than 600 times the interaction length of 10^{15} eV photons with the 3 K microwave background radiation. Subsequently UHE γ -rays above 10^{15} eV have been observed from another extragalactic source, LMC X-4, in phase with its 1.408^h orbital period (Protheroe and Clay, 1985). This object is in the Large Magellanic Cloud at a distance of ~ 50 kpc, i.e. about 5 interaction lengths away. With this recent extension of extragalactic astronomy to photon energies in excess of 10^{14} eV, it is therefore timely to re-examine the extragalactic propagation of γ -rays in detail. Here I report the result of Monte Carlo electron-photon cascade calculations for propagation of γ -rays through regions of extragalactic space containing no magnetic field. These calculations then provide upper limits to the expected flux from extragalactic sources.

2. **THE SIMULATION.** Since we are interested primarily in γ -rays in the 10^{14} – 10^{17} eV energy range, I have only considered interactions of electrons and photons with the 3 K microwave background radiation. As mentioned above, to obtain an upper limit to the expected γ -ray flux from sources, I have assumed that the intergalactic field is so low that it can be ignored (i.e. $\ll 10^{-7}$ gauss). Interactions with photons of the near-infrared background radiation (Matsumoto *et al.*, 1983; de Bernadis *et al.*, 1985) are not considered here although these will have important implications for γ -rays below 10^{14} eV if the near-infrared background radiation is universal (this topic is discussed later). Interaction lengths of electrons and photons in the microwave background radiation at a temperature of 2.96 K (Narlikar, 1982) have been calculated and are given in Fig. 1. For photon-photon interactions, this is based on the results of Gould and Schreder (1967b) together with the correction by Gould (1983) and, for IC interactions, on an integration over the full Klein-Nishina cross sections (Jauch and Rohrlich, 1976).

Full Monte Carlo simulations of the electron-photon cascade involving photon-photon pair production interactions and IC interactions of electrons with the 3 K background were then performed using the exact cross sections. Full details of the calculation will be published elsewhere (Protheroe, 1985).

To run the Monte Carlo simulation program over very large distances (e.g. greater than ~200 kpc) would take excessive computer time. To get around this problem, the simulation was run over a smaller distance, Δx , in order to obtain matrices $[A(\Delta x)]$, $[B(\Delta x)]$, $[C(\Delta x)]$, and $[D(\Delta x)]$ which describe the evolution of arbitrary electron and photon spectra after propagation through distance Δx of extragalactic space. The energy spectra of photons and electrons at distance x from the source, $F_\gamma(E, x)$ and $F_e(E, x)$, are represented by column vectors $f^\gamma(x)$ and $f^e(x)$ whose elements are defined by

$$f^\gamma_k(x) = \int_{E_k}^{\alpha E_k} F_\gamma(E, x) dE \quad (1)$$

and

$$f^e_k(x) = \int_{E_k}^{\alpha E_k} F_e(E, x) dE \quad (2)$$

where $\alpha = E_{k+1}/E_k$. The spectra of photons and electrons after traversing a further distance Δx are then obtained from

$$f^\gamma(x+\Delta x) = [A(\Delta x)]f^\gamma(x) + [B(\Delta x)]f^e(x) \quad (3)$$

$$f^e(x+\Delta x) = [C(\Delta x)]f^\gamma(x) + [D(\Delta x)]f^e(x) \quad (4)$$

where $A(\Delta x)_{kj}$ is the mean number of photons produced with energies in the range E_k to αE_k as a result of an electron-photon cascade through distance Δx initiated by an "average" photon with an energy in the range E_j to αE_j . The elements of $[B(\Delta x)]$, $[C(\Delta x)]$, and $[D(\Delta x)]$ are defined in a similar way for the appropriate primary and secondary species.

In the present simulation, $E_1 = 100$ GeV, $\alpha = 10^{0.1}$, and 90 energy bins were used. For each primary energy, e.g. bin k , 400 primary γ -rays were sampled from an E^{-2} differential spectrum in the range E_k to αE_k and for each primary γ -ray a full Monte Carlo calculation was carried out over a distance $\Delta x = 50$ kpc. The secondary photons and electrons reaching distance Δx having energies in bin j were then divided by 400 to obtain $A(\Delta x)_{kj}$ and $B(\Delta x)_{kj}$ respectively. An identical method was adopted for primary electrons to obtain $C(\Delta x)_{kj}$ and $D(\Delta x)_{kj}$.

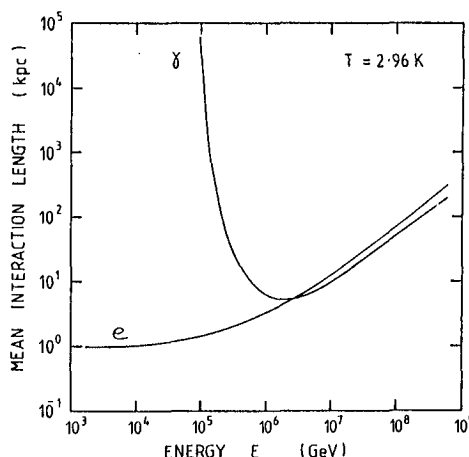


Fig. 1. Mean interaction length of electrons and γ -rays in a 2.96 K black body radiation field.

3. RESULTS. The differential and integral energy spectra of γ -rays

expected at various distances ranging from 50 kpc to 6.4 Mpc from a source which emits an E^{-2} differential γ -ray spectrum are shown in Fig. 2 for two extreme assumptions about the intergalactic magnetic field.

The spectra corresponding to a distance of 6.4 Mpc would apply to Cen A. Clearly, if cascading takes place then γ -rays at 10^{14} – 10^{16} eV energies could reach Earth. To see what primary γ -ray energies (i.e. at the source) would produce γ -rays which would be observed in this energy range, the spectrum expected from a monoenergetic source of γ -rays at a distance of 6.4 Mpc has been calculated for various primary energies in the range 10^{14} – 10^{16} eV. The result is shown in Fig. 3 and indicates that γ -rays at 10^{14} – 10^{16} eV energies from a source at 6.4 Mpc distant would be due to primary γ -rays of energy $\sim 10^{14}$ eV.

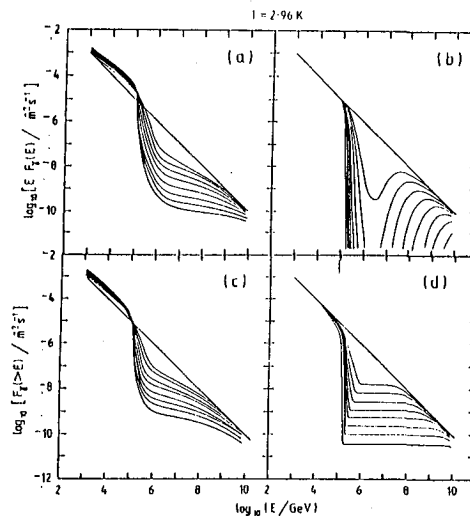


Fig. 2. Expected differential (a and b) and integral (c and d) energy spectra of γ -rays at distances of 0, 50, 100, 200, 400, 800 kpc, 1.6, 3.2 and 6.4 Mpc from a source emitting an E^{-2} differential spectrum of γ -rays (the curves are not labelled but may be identified as in each case the flux at 10^6 GeV decreases monotonically as distance is increased). (a) and (c) are for zero magnetic field conditions in which cascading is unimpeded; (b) and (d) are for high field conditions. (N.B. curves do not take account of the inverse-square law).

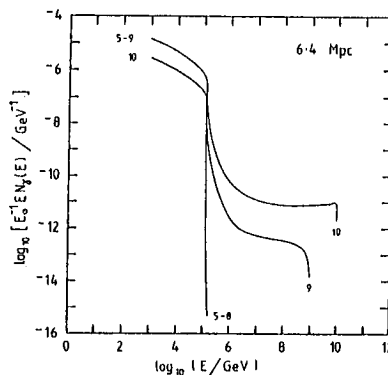


Fig. 3. Average differential energy spectrum expected at Earth from a source at a distance of 6.4 Mpc emitting monoenergetic γ -rays. The curves do not take account of the inverse-square law and are normalised to one primary γ -ray and are divided by the primary energy, E_0 . (Numbers attached to the curves are $\log_{10}(E_0/\text{GeV})$).

4. DISCUSSION. Sources in the Magellanic Clouds, or in other galaxies of the local group, should have energy spectra which display an absorption feature at 10^{14} – 10^{16} eV energies together with a peak just below 10^{14} eV. The peak just below 10^{14} eV could make observing in this energy range quite attractive. The shape of the spectrum contains information on both the magnetic field and the column density of microwave photons, i.e. the relative heights of the peak and dip will depend on the magnetic field and distance to the source. If the energy spectra could be accurately mapped out, the depth of the absorption feature together with the height of the peak could then possibly be used to measure both the distance to the source and the average strength of the magnetic field along the line of sight to the source. It is thus possible, in principle at least, to obtain information about the radiation content and magnetic field in intergalactic space within the local group by studying the energy spectra of γ -ray

sources.

It is clear from Fig. 2(a) that it is possible for $\sim 10^{12}$ eV γ -rays to reach Earth from Cen A provided the intergalactic magnetic field is low enough. Even then, however, the flux is down by $\sim 10^{-2}$ of that we would receive if the cosmic microwave background radiation were not present. For the observed excess of air showers from the direction of Cen A (Clay et al., 1984) to be due to γ -rays from Cen A, either the microwave background radiation must not be universal or Cen A must be much more luminous at $\sim 10^{12}$ eV energies than at lower energies. It would have to have a luminosity per decade at $\sim 10^{12}$ eV comparable to that of 3C273 at lower energies. A third possibility is that a foreground source in our Galaxy is being observed. Clearly, additional observations of Cen A are urgently required.

There has recently been an observation (Matsumoto et al., 1983) of a dilute near-infrared background radiation. If the near-infrared photon energy density is as high as suggested and is universal then sources more distant than any in the local group of galaxies would show energy spectra strongly attenuated below 10^{14} down to $\sim 10^{12}$ eV (see also Rana and Wolfendale, 1984). A peak in the energy spectrum would not then be apparent at just below 10^{14} eV for these sources but might occur instead at just below 10^{12} eV. This could be the explanation for the somewhat higher than expected (on the basis of extrapolation from hard X-ray energies through 100 MeV upper limits; see e.g. the survey by Baitty et al., 1981) γ -ray flux from Cen A observed by Grindlay et al. (1975) at these energies.

Acknowledgements This work has been supported by the A.R.G.S.

REFERENCES

- Allcock MC and Wdowczyk J, 1972: *Nuovo Cim.*, 9B, 315-320
 Baitty WA, et al., 1981: *Ap. J.*, 244, 429-435
 de Bernardis P, et al., 1985: *Ap. J.*, 288, 29-31.
 Bonnetto SA, 1971: *Nuovo Cim. Lett.*, 1, 677-680
 Bonnetto SA and Lucchin F, 1971: *Nuovo Cim. Lett.*, 2, 1299-1304
 Brown RW, et al., 1973: *Ap. Lett.*, 14, 203-205
 Clay RW, et al., 1984: *Aust. J. Phys.*, 37, 91-95
 Gould RJ, 1983: *Ap. J. (Lett.)*, 274, L23-L25
 Gould RJ and Rephaeli Y, 1978 *Ap. J.*, 225, 318-324
 Gould RJ and Schreder G, 1966: *Phys. Rev. Lett.*, 16, 252-254
 Gould RJ and Schreder G, 1967a: *Phys. Rev.*, 155, 1408-1411
 Gould RJ and Schreder G, 1967b: *Phys. Rev.*, 155, 1404-1407
 Grindlay J, et al., 1975: *Ap. J.*, 201, 82-89
 Jauch JM and Rohrlich F, 1976: "The Theory of Photons and Electrons" (Berlin: Springer-Verlag) p. 234
 Jelley JV, 1966: *Phys. Rev. Lett.*, 16, 479-481
 Matsumoto T, et al., 1983: preprint
 Narlikar J, 1982: "Violent Phenomena in the Universe" (Oxford: Oxford University Press) p. 155
 Protheroe RJ, 1985: in preparation
 Protheroe RJ and Clay RW, 1985: *Nature*, in press; see also these proceedings paper OG 2.6-10
 Rana NC and Wolfendale AW, 1984: *Vistas in Astron.*, 199-216
 Stecker FW, 1973: *Ap. Space Sci.*, 20, 47-57
 Wdowczyk J, et al., 1972: *J. Phys. A: Gen. Phys.*, 5, 1419-1432